

# Measuring $D_A$ and $H$ at $z=0.35$ from the SDSS DR7 LRGs using Baryon Acoustic Oscillations



arXiv: 1206.6732

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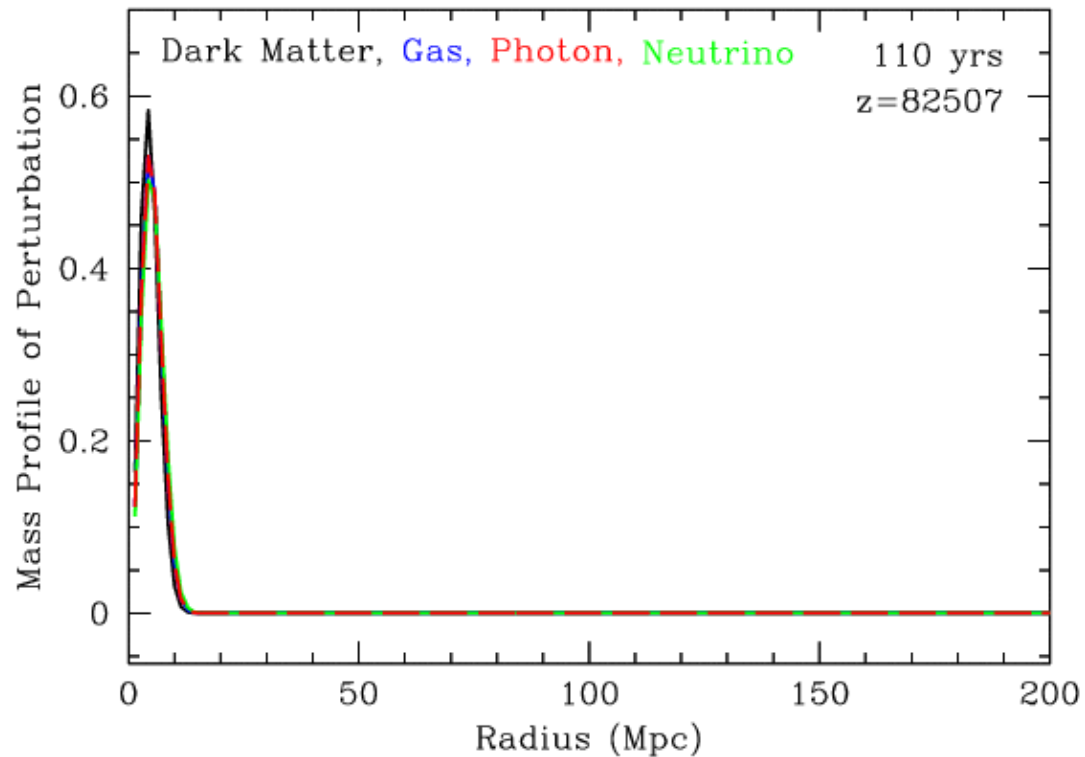
Workshop on Cosmic Acceleration

Carnegie Mellon University

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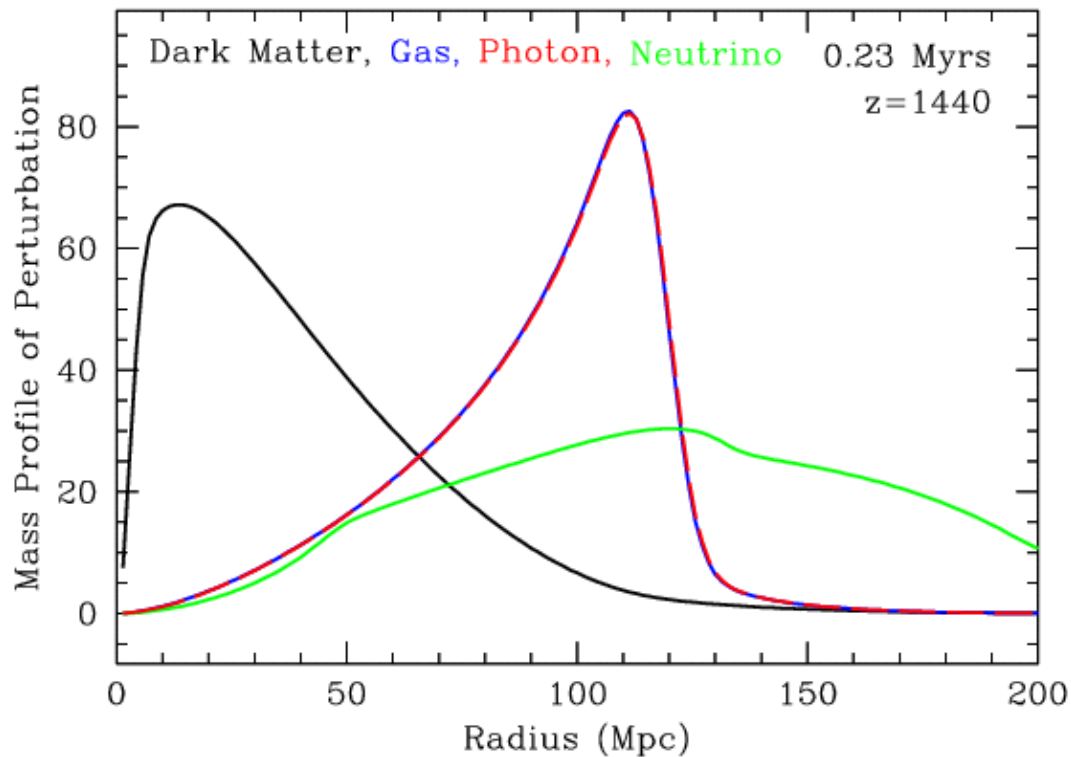
# What are BAO?



[http://cmb.as.arizona.edu/~eisenste/acousticpeak/acoustic\\_physics.html](http://cmb.as.arizona.edu/~eisenste/acousticpeak/acoustic_physics.html)

We start with a tiny overdensity in the primordial universe...

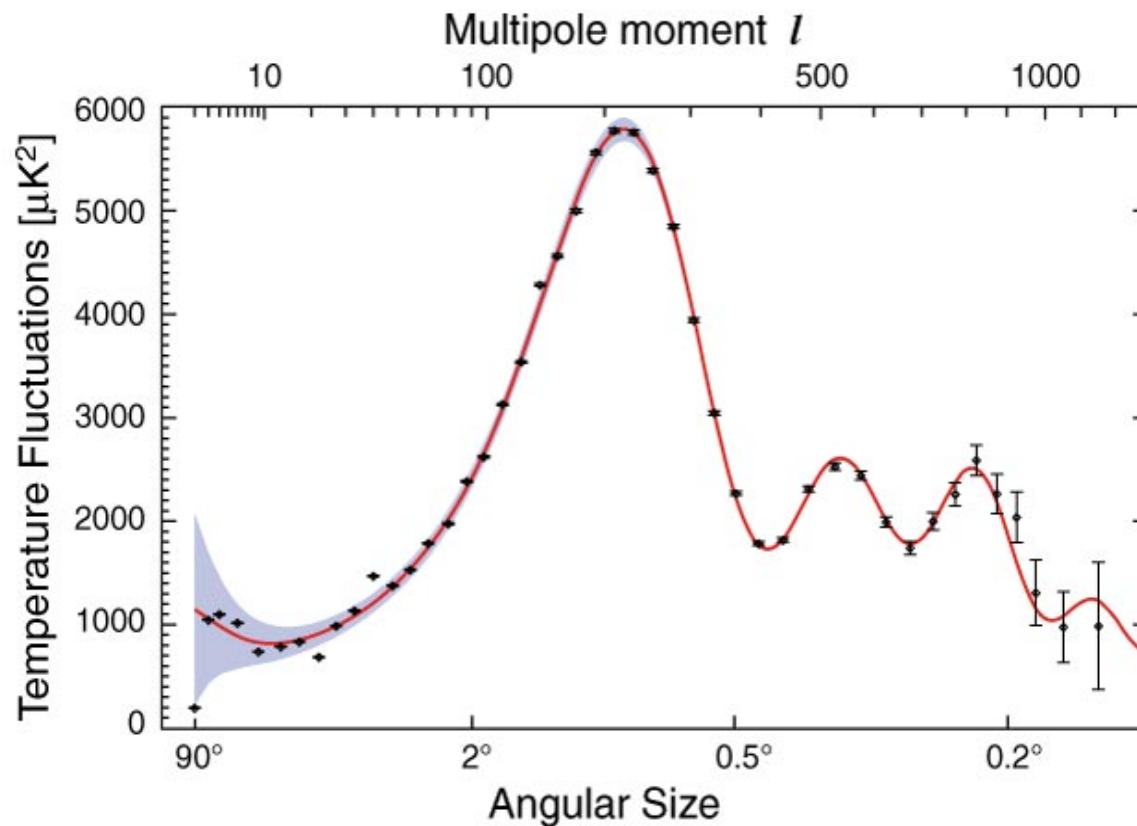
# What are BAO?



Photons and baryons push out in a spherical sound wave.

Dark matter tries to pull material back in.

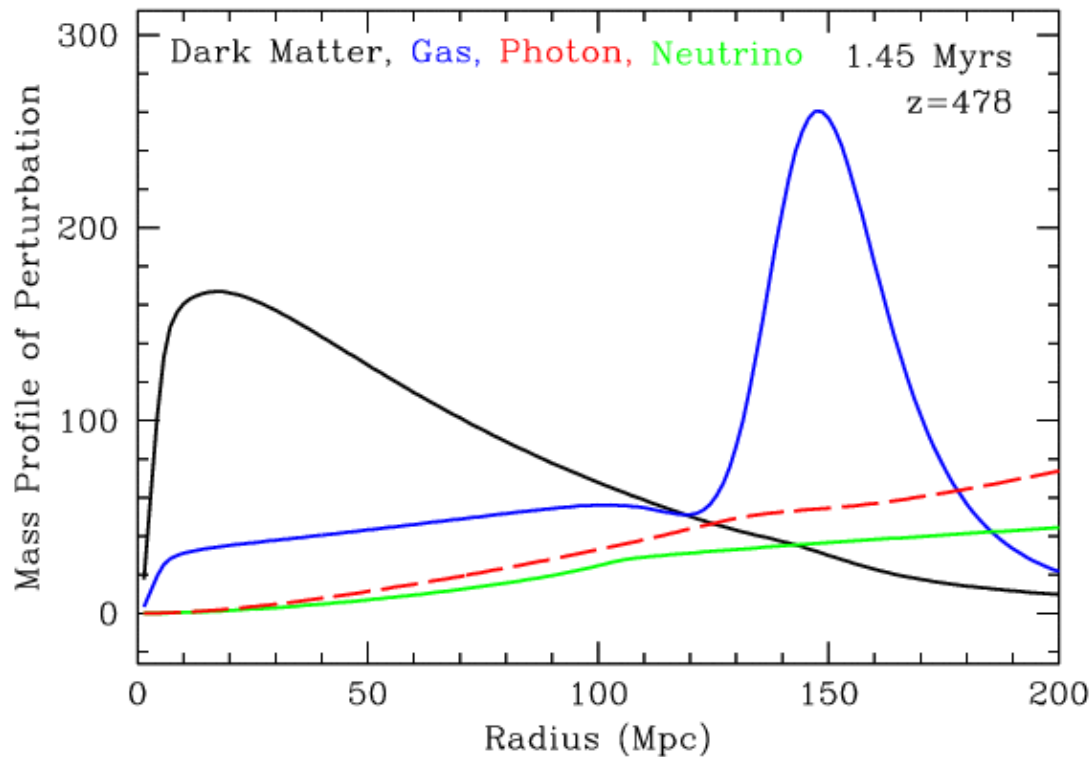
# What are BAO?



Standing waves form in the fluid.

At recombination, their phases are imprinted on the photon and baryon distributions; these are the **baryon acoustic oscillations**.

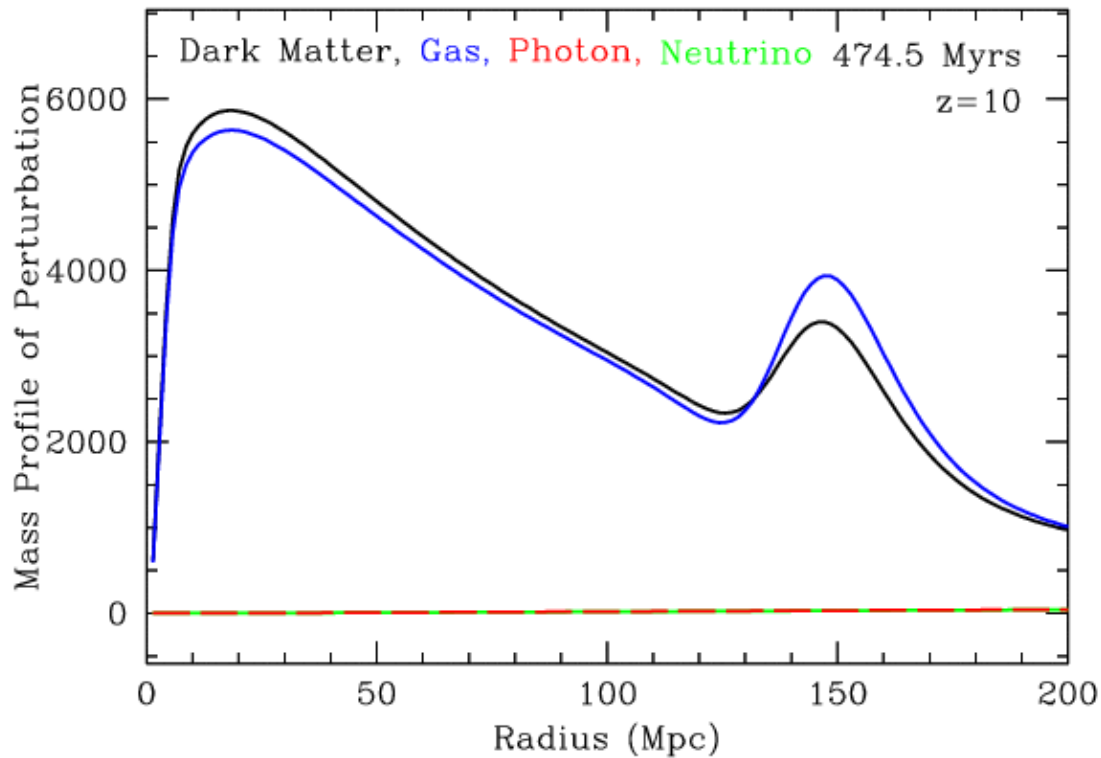
# What are BAO?



The sound waves travel about **150 comoving Mpc** prior to recombination.

Baryons deposited at this distance.

# What are BAO?



Gravity equilibrates the distributions of dark matter and baryons.

A tiny excess is left at 150 Mpc; this is known as the **acoustic peak**.

# The BAO as a Standard Ruler

- The position of the acoustic peak marks a characteristic scale that is about 150 comoving Mpc.
- This scale is known as the **acoustic scale** or the **sound horizon**.
- Its magnitude only depends on 2 factors:
  - Time of matter radiation equality (depends on  $\Omega_m h^2$ ).
  - Value of  $\Omega_b h^2$ .
- The sound horizon has been measured to 1.1% accuracy by WMAP7 → it can be used as a very accurate **standard ruler**.

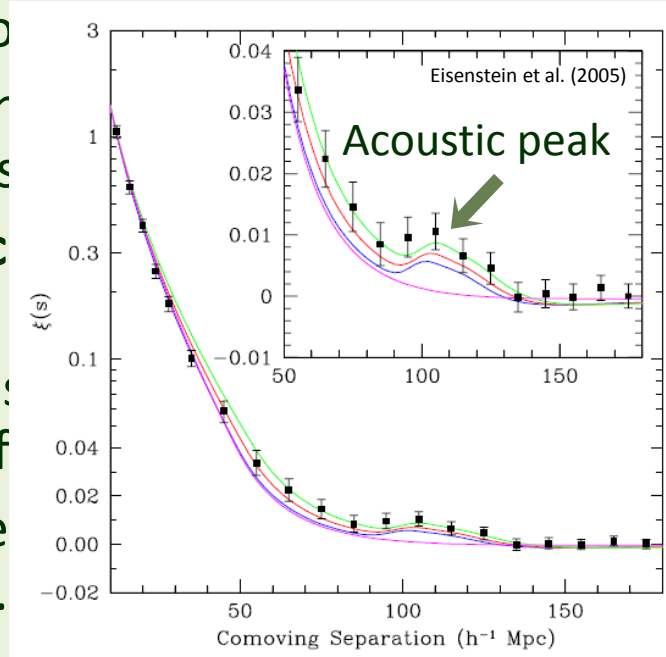
# The BAO as a Standard Ruler

- Start with a set of galaxy observations covering a large volume.
- Compute the spherically-averaged correlation function assuming some fiducial cosmology.
- **If the assumed cosmology is wrong, the BAO will appear slightly shifted.**
- Parameterize this shift in a model and fit model to data to measure the shift.
- Repeat at different redshifts => spherically-averaged distance-redshift relation.



# The BAO as a Standard Ruler

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# The BAO as a Standard Ruler

- Start with a set of galaxy observations covering a large volume.
- Compute the spherically-averaged correlation function assuming some fiducial cosmology.

- **If the assumed cosmology is slightly shifted.**

$$\alpha = \frac{D_V(z)/r_s}{D_{V,f}(z)/r_{s,f}}$$

will appear

- Parameterize this shift as  $\alpha$  and measure the shift.

$$= \left[ \frac{D_A^2(z) H_f(z)}{D_{A,f}^2(z) H(z)} \right]^{1/3} \frac{r_{s,f}}{r_s}$$

fit to data to

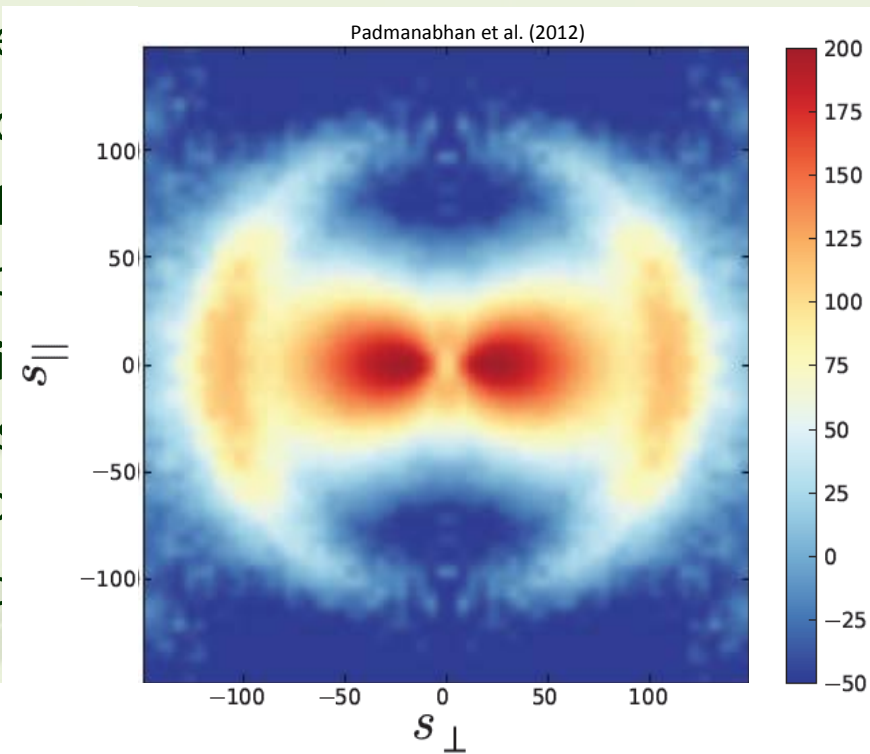
- Repeat at different redshifts => spherically-averaged distance-redshift relation.

# Anisotropic BAO

- Anisotropies in the BAO can help us separate  $D_A$  and  $H$ .
- Anisotropic clustering arises from 2 sources:
  - **Redshift space distortions:** Cause changes to the broadband, does not affect the BAO in particular.
  - **Assuming the wrong cosmology:** Shifts the BAO such that it appears in slightly different positions along different directions.
- The second of these can be used to separately constrain  $D_A(z)$  and  $H(z)$ , allowing us to directly probe the cosmic expansion history.

# Anisotropic BAO

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# Anisotropic BAO

- Anisotropic clustering introduces power into the higher order even multipoles of the correlation function.
- Hence, to measure the anisotropy in the BAO, we can use the  $l=2$  quadrupole (we will still assume that  $l>2$  moments are 0).
- We fit models of the monopole and quadrupole to the data; these models contain terms that parameterize the BAO shift:

$$1 + \varepsilon = \sqrt[3]{\frac{H_f(z)D_{A,f}(z)}{H(z)D_A(z)}}$$

- We then simultaneously fit the monopole and quadrupole for  $\alpha$  and  $\varepsilon \rightarrow D_A$  and  $H$ .

# Anisotropic BAO

- Anisotropic clustering introduces power into the higher order even multipoles of the correlation function.

- Hence, to measure  $\alpha$ , we can use the  $l=2$  quadrupole (where the odd moments are 0).

$$\alpha = \frac{D_V(z)/r_s}{D_{V,f}(z)/r_{s,f}}$$

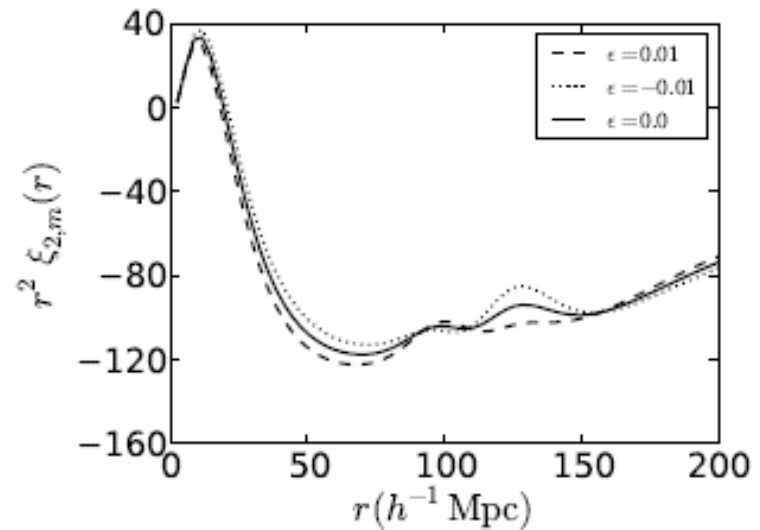
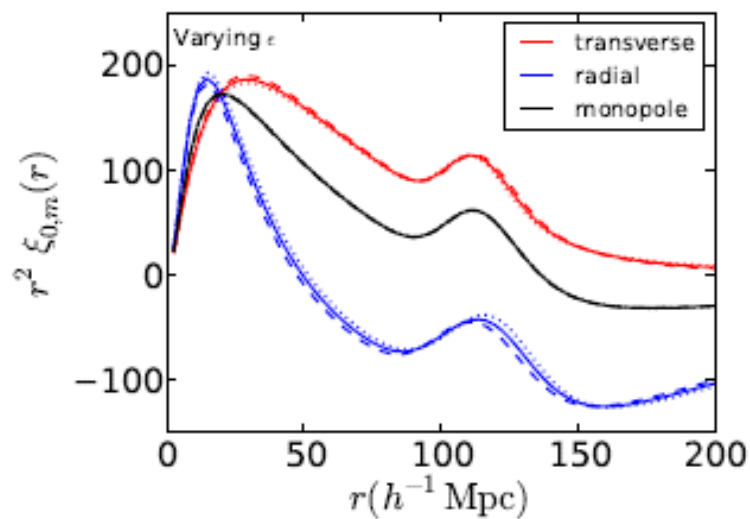
$$= \left[ \frac{D_A^2(z)}{D_{A,f}^2(z)} \frac{H_f(z)}{H(z)} \right]^{1/3} \frac{r_{s,f}}{r_s}$$

- We fit models of  $\alpha$  to the data; we the BAO shift:

$$1 + \varepsilon = \sqrt[3]{\frac{H_f(z)D_{A,f}(z)}{H(z)D_A(z)}}$$



- We then simultaneously fit the monopole and quadrupole for  $\alpha$  and  $\varepsilon \rightarrow D_A$  and  $H$ .

# The Quadrupole



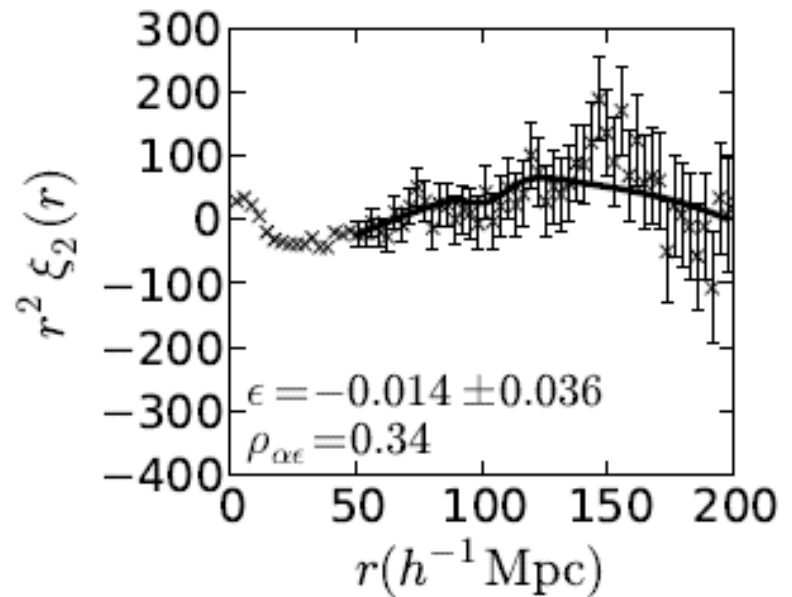
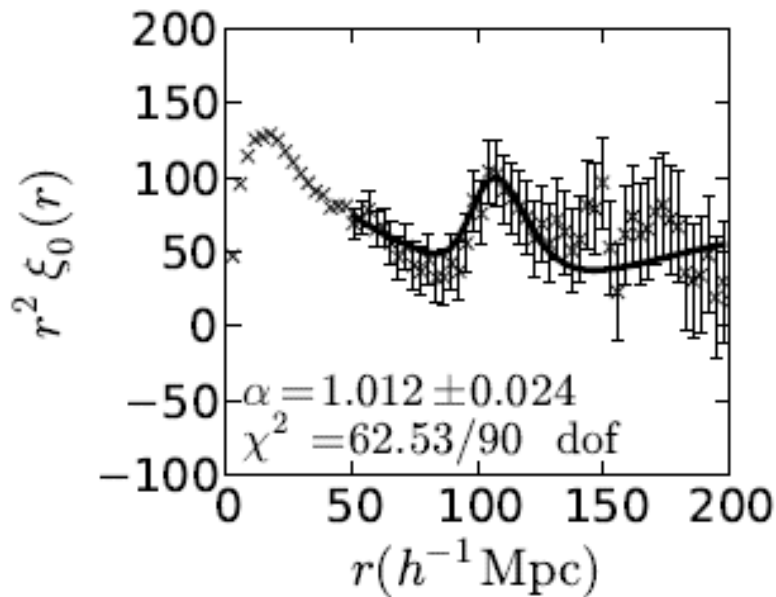


# Our Data

- We use data from the Sloan Digital Sky Survey (SDSS).
  - Data taken by 2.5m telescope at Apache Point Observatory.
  - SDSS is currently in its 3rd generation; we use the final data release from the 2nd generation (SDSS-II DR7).
  - **DR7 LRG sample** (2 spectrographs, 600 fibers per):
    - $0.16 < z < 0.47$  (median  $z=0.35$ ), about 7000 deg<sup>2</sup> sky coverage,  $1 \times 10^{-4} h^3/\text{Mpc}^3$  number density
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# DR7 Results



$H(z) = 84.4 \pm 7.1 \text{ km/s/Mpc}$  (8.4% measurement)

$D_A(z) = 1050 \pm 38 \text{ Mpc}$  (3.6% measurement)



# Conclusions

- We present methods for anisotropic BAO analysis and use it to separately constrain  $D_A(z)$  and  $H(z)$  using SDSS DR7.
  - We obtain an 8.4% measurement of  $H(z)$  and a 3.6% measurement of  $D_A(z)$ .
  - Future BAO surveys will be able to improve the precision of these measurements; the methods presented here will be directly applicable to future analyses.
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